



Early Journal Content on JSTOR, Free to Anyone in the World

This article is one of nearly 500,000 scholarly works digitized and made freely available to everyone in the world by JSTOR.

Known as the Early Journal Content, this set of works include research articles, news, letters, and other writings published in more than 200 of the oldest leading academic journals. The works date from the mid-seventeenth to the early twentieth centuries.

We encourage people to read and share the Early Journal Content openly and to tell others that this resource exists. People may post this content online or redistribute in any way for non-commercial purposes.

Read more about Early Journal Content at <http://about.jstor.org/participate-jstor/individuals/early-journal-content>.

JSTOR is a digital library of academic journals, books, and primary source objects. JSTOR helps people discover, use, and build upon a wide range of content through a powerful research and teaching platform, and preserves this content for future generations. JSTOR is part of ITHAKA, a not-for-profit organization that also includes Ithaka S+R and Portico. For more information about JSTOR, please contact support@jstor.org.

β 733 = 85 *Pegasi*.—The angular separation in this pair is still great enough to make the little companion easily visible under favorable conditions. Three recent measures give:—

1903.54 262°.0 0".76

The computed position from BURNHAM's orbit for the same date is: 267°.9, 0".72.

September 10, 1903.

R. G. AITKEN.

PERSONAL SCALE.

Dr. SCHLESINGER's note on this subject in No. 91 of these *Publications* is of interest to all who are engaged in work involving the use of micrometer-screws. For the sake of comparison, I have thought it might be worth while to give here the results of a similar count I have made of the number of times each tenth was recorded in 1,200 recent readings of the micrometers attached to the 36-inch and 12-inch telescopes, together with the results of a similar count of 1,000 earlier readings. As the milled heads of the micrometer-screws are divided into one hundred parts, the tenths of divisions here considered have the value of thousandths of one revolution. The personal scale, therefore, does not affect the accuracy of the micrometer measures. In the following table all of the percentages would be 10.0 if there were no personal scale.

Tenths.	1200	1000
.0	20.2	20.3
.1	8.2	8.9
.2	16.7	15.7
.3	8.0	8.4
.4	6.8	7.8
.5	9.2	11.0
.6	5.9	6.7
.7	5.3	3.0
.8	15.3	14.6
.9	4.4	3.6

It is apparent that the two counts give practically the same error-curve, and show that while 0.5 was read as often as it occurred, 0.1, and especially 0.9, were slighted in favor of 0.0; and 0.3, 0.4, and 0.6, 0.7 in favor of 0.2 and 0.8 respectively.

September 8, 1903.

R. G. AITKEN.

DISCOVERY OF A NEW COMPANION TO O Σ 499.

In the course of my survey of the stars in the zone $+56^\circ$ to $+60^\circ$, I turned the 36-inch telescope upon O Σ 499 on the night of June 11, 1903, and found that the fainter star of this wide pair was a close double.

Three nights' measures of the new pair give:—

1903.54 $146^\circ.6$ $0''.42$ 9.0-10.8 magnitudes.

The wide pair is probably fixed, having shown no motion since discovery in 1847.

September 10, 1903.

R. G. AITKEN.

REDISCOVERY OF COMET 1889 V, 1896 VI (BROOKS).

The first observation of BROOKS's Comet during the present apparition was secured with the 36-inch telescope on August 18th, $12^h 17^m 23^s$ Mt. Hamilton M. T., at which time it was found in R. A. $21^h 2^m 51^s.30$; Decl. $-27^\circ 4' 18''.8$. The position predicted for this date by Professor NEUGEBAUER differs from the observed place by but little more than $5'$. The comet is therefore following the predicted path very closely.

This comet is of unusual interest for several reasons. At the time of its first appearance, in 1889, it was attended by several companions, one of which at times rivaled its primary in brightness. It is also interesting because of its possible identity with the lost LEXELL's comet, and because it is the only comet that has been observed at two consecutive oppositions, Professor BARNARD having secured five observations of it between November 22, 1890, and January 13, 1891,—a year and a half after its discovery.

On its second return to the Earth it was first observed by M. JAVELLE, at Nice, on June 20, 1896, and was followed by observers at fourteen stations until February 25, 1897, when Professor HUSSEY secured the last observation, with the 36-inch telescope.

Professor BAUSCHINGER determined the definitive elements from the observations in 1889-91, and later revised his work with the aid of the observations made in 1896-97. These elements, with corrections applied for perturbations, afforded the basis for the ephemeris which Professor NEUGEBAUER computed for the present apparition.